Table 1. Recent Distance Moduli to the LMC

No.	Method	Object	Author	m-M
1	Baade-Wesselink	Cepheids	Gieren et al. 2000	18.42 ± 0.10
2		Cepheids	Carretta et al. 2000a	18.55 ± 0.10
3		Cepheid	Gieren, Fouque, & Gomez 1998	18.46 ± 0.02
4		Cepheid	Di Benedetto 1997	18.58 ± 0.024
5		RR Lyraes	Carretta et al. 2000b	18.52 ± 0.20
6		RR Lyraes	Feast 1997	18.53 ± 0.04
7		RR Lyraes	McNamara 1997	18.54 ± 0.10
8	Double-mode	RR Lyraes	Kovacs 2000	18.52 ± 0.21
9		RR Lyraes	Alcock et al. 1997	18.48 ± 0.19
10	Eclipsing binaries	EROS 1044	Ribas et al. 2002b	18.38± 0.08
11		EROS 1044	Ribas et al. 2002a	18.25 ± 0.25
12		EROS 1044	Maloney et al. 2001	18.25 ± 0.25
13		HV 982	Salaris & Groenewegen 2002	18.43 ± 0.24
14		HV 982	Fitzpatrick et al. 2002	18.52 ± 0.06
15		HV 982	Fitzpatrick 2000	18.31 ± 0.09
16		HV 2274	Salaris & Groenewegen 2002	18.42 ± 0.24
17		HV 2274	Fitzpatrick et al. 2002	18.31 ± 0.1
18		HV 2274	Groenewegen & Salaris 2001	18.46 ± 0.07
19		HV 2274	Nelson et al. 2000	18.40 ± 0.07
20		HV 2274	Guinan et al. 1998a	18.42 ± 0.07
21		HV 2274	Guinan et al. 1998b	18.30 ± 0.07
22		HV 2274	Udalski 1998a	18.22 ± 0.13
23		HV2274	Guinan et al. 1997	18.54 ± 0.08
24	Globular Cluster Dyn. mods		Chaboyer et al. 1998	18.50 ± 0.11
25	High Amplitude δ Scuti	δ Scuti	McNamara 2001	18.66 ± 0.08
26	Long Period Variables	Mira	Feast et al. 2002	18.60 ± 0.10
27		Mira	Whitelock & Feast 2000	18.64 ± 0.17
28		Mira	Van Leeuwen et al. 1997	18.54 ± 0.18
29		Ca Stars	Bergeat, Knapik, & Rutily 1998	18.50 ± 0.17
30	M Stars Luminosity		Schmidt-Kaler & Oestreicher 1998	18.34 ± 0.09
31	Main Sequence fitting	NGC 1866	Walker et al. 2001	18.33 ± 0.05
32		NGC 1466 & 2257	Johnson et. al 1999	18.46 ± 0.09
32		& Hodge 11		
33		Cepheids	Carretta et al. 2000a	18.55 ± 0.04 ± 0.04
34		Cepheids	Laney & Stobie 1994	18.49 ± 0.04 ± 0.04
35	Masers	NGC 4258	Caputo et al. 2002	18.5 ± 0.10
36		NGC 4258	Caputo et al. 2002	18.15 to 18.45 [k]
37		NGC 4258	Caputo et al. 2002	18.35 to 18.6 [l]
38		NGC 4258	Newman et al. 2001	18.31 ± 0.11 ± 0.17
39	Mean V magnitude	LMC RR Lyraess	McNamara 2001	18.61 ± 0.04
40	Modelling Li-rich Ca stars		Ventura, D'Antona,	18.70 ± 0.25
40			& Mazzitelli 1999	
41	Pulsation modelling	Cepheids	Keller 2002	18.55 ± 0.02
42		Cepheids	Bono et. al 2002	18.54 ± 0.09 [m]
43		Cepheids	Bono et. al 2002	18.42 ± 0.17 [pm]
44		Cepheids	Bono et. al 2002	18.51 ± 0.14 [qm]
45		Cepheids	Bono et. al 2002	18.48 ± 0.14 [n]
46		Cepheids	Bono et. al 2002	18.28 ± 0.16 [p]
47		Cepheids	Bono et. al 2002	18.45 ± 0.35 [q]

Table 1—Continued

No.	Method	Object	Author	m-M
48		Cepheids	Bono, Catellani & Marconi 2002	18.48 to 1858
49		Cepheids	Wood 1998	18.54 ± 0.08
50	Planetary Nebulae Luminosity	M31	Walker 1999	18.50 ± 0.18
51	Red Clump		Alves et. al 2004	18.493 ± 0.033r $\pm$ 0.03s
52			Clementini et. al 2001	18.47 ± 0.1
53			Popowski 2001	18.33 ± 0.07 [a]
54			Girardi & Salaris 2001	18.55 ± 0.05
55			Sakai, Zaritsky & Kennicutt 2000	18.29 ± 0.03
56			Popowski 2000	18.27 ± 0.06 [b]
57			Stanek et al. 2000	18.24 ± 0.08
58			Romaniello et al. 2000	$18.59 \pm 0.04 \pm 0.08$
59			Udalski 2000a	18.24 ± 0.08
60			Twarog, Anthony-Twarog	18.42 ± 0.16
60			& Bricker 1999	
61			Stanek, Zaritsky, & Harris 1998	18.065 ± 0.12
62			Udalski et al. 1998b	18.08 ± 0.15
63			Udalski 1998a	18.09 ± 0.16
64			Udalski 1998b	18.18 ± 0.06
65			Cole 1998	18.36 ± 0.17
56			Girardi et al. 1998	18.28 ± 0.14
67			Beaulieu & Sackett 1998	18.3
68	Red Clump & RR Lyraes		Popowski 2001	18.24 ± 0.08 to 18.44 ± 0.07
69	SN 1987A		Carretta et al. 2000	18.58 ± 0.05
70			Romaniello et al. 2000	18.55 ± 0.05
71			Walker 1999	$18.55 \pm 0.07 \pm 0.16$
72			Gould & Uza 1998	18.37 ± 0.04
73			Panagia, Gilmozzi,	18.58 ± 0.08
73			& Kirchner 1998	
74			Lundqvist & Sonneborn 1998	18.67 ± 0.05
75	Statistical parallaxes	RR Lyraes	Carretta et al. 2000a	18.38 ± 0.12
76		RR Lyraes	Popowski & Gould 1999	18.33 ± 0.08 [a]
77		RR Lyraes	Popowski & Gould 1999	18.23 ± 0.08 [b]
78		RR Lyraes	Popowski & Gould 1998	18.07 ± 0.15 [a]
79		RR Lyraes	Popowski & Gould 1998	18.31 ± 0.14 [b]
80		RR Lyraes	Gould & Popowski 1998	18.24 ± 0.14
81		RR Lyraes	Fernley et al. 1998	18.26 ± 0.15
82		RR Lyraes	Layden et. al 1996	18.28 ± 0.03
83	Subdwarf fitting		Carreta et al. 2000a	18.64 ± 0.12
84			Reid 1998	18.79 ± 0.17
85			Reid 1997	18.65 ± 0.12
86	Tip of the Red Giant Branch		Sakai et al. 2000	$18.59 \pm 0.09 \pm 0.16$
87			Romaniello et al. 2000	$18.69 \pm 0.25 \pm 0.06$
88			Cioni et al 2000	$18.55 \pm 0.04 \pm 0.08$
89	Trigonometric parallax	RR Lyrae	Benedict et al. 2002	18.53 ± 0.10 [c]
90		RR Lyrae	Benedict et al. 2002	18.38 ± 0.10 [d]
91		RR Lyrae	Luri et al. 1998	18.37 ± 0.23
92		RR Lyrae	McNamara 1997	18.57 ± 0.03
93		Cepheids	Benedict et al. 2002	18.50 ± 0.13 [b]
94		Cepheids	Benedict et al. 2002	18.58 ± 0.15 [j]

Table 1—Continued

No.	Method	Object	Author	m-M
95		Cepheids	Feast et al. 2002	18.6 ± 0.10
96		Cepheids	Groenewegen & Oudmaijer 2000	18.45 ± 0.18 to 18.86 ± 0.12
97		Cepheids	Groenewegen & Oudmaijer 2000	18.60 ± 0.11 [e]
98		Cepheids	Groenewegen & Oudmaijer 2000	18.52 ± 0.18 [f]
99		Cepheids	Groenewegen 2000	18.60 ± 0.10 [g]
100		Cepheids	Groenewegen 2000	18.55 ± 0.17 [h]
101		Cepheids	Groenewegen & Salaris 1999	18.61 ± 0.28
102		Cepheids	Feast 1999	18.68 ± 0.22 [i]
103		Cepheids	Oudmaijer, Groenewegen,	18.56 ± 0.08
103			& Schrijver 1998	
104		Cepheids	Madore & Freedman 1998	18.44 ± 0.35 to 18.57 ± 0.11
105		Cepheids	Luri et al. 1998	18.29 ± 0.17
106		Cepheids	Feast & Catchpole 1997	18.70 ± 0.10
107		Cepheids	Paturel et al. 1997	18.72 ± 0.05
108		HB	Carretta et al. 2000a	18.49 ± 0.11
109		HB	Gratton 1998	18.49 ± 0.11
110		HB	Koen & Laney 1998	18.49 ± 0.12
111	White Dwarf cooling sequence		Carretta et al. 2000a	18.40 ± 0.15

[a] Walker (1992) Photometry

[b] Udalski (1999) Photometry

[c] Caretta (2000) $<V>$ = 19.11 (+0.03, $<$[FeH]$>$ correction)=19.14

[d] Udalski (1999) $<V>$ = 18.94 (+0.05, $<$[FeH]$>$ correction) = 18.99

[e] Based on the *PL*-relation in V & I & the Wesenheit-index

[f] Based on the *PL*-relation in K

[g] Based on the *PL*-relation in W

[h] Based on the *PL*-relation in K

[i] Data from Koen & Laney 1998

[j] Highly uncertain metallicity correction (Freeman et al. 2001)

[k] empirical metallicity correction adopted

[l] theoretical metallicity correciton adopted

[m] using predicted FO *PL*-relations

[n] using empiracal FO *PL*-relations

[o] W solution

[p] I_0 solution

[q] K_0 solution

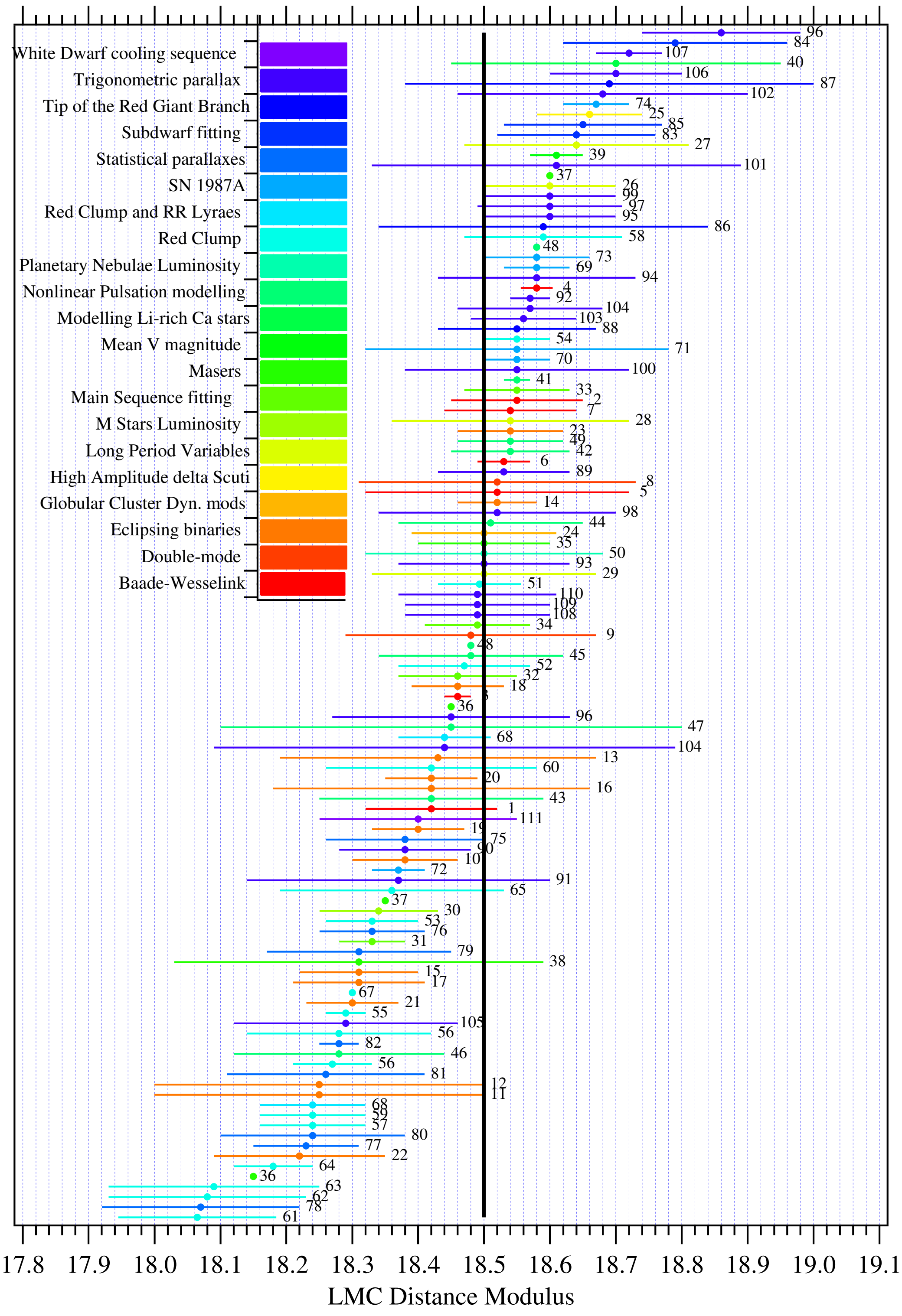

Fig. 1.— Recent determinations of the distance modulus of the Large Magellanic Cloud, an expansion of the plot found in Gibson (1999). Colors represent the various methods listed in column 2 of Table 1, while the numbers refer to the individual investigations (column 1). The thick vertical line denotes the distance modulus adopted by the HST Distance Scale Key Project (Freedman *et al.* 2001) and the Type Ia Supernovae Calibration Team (().).

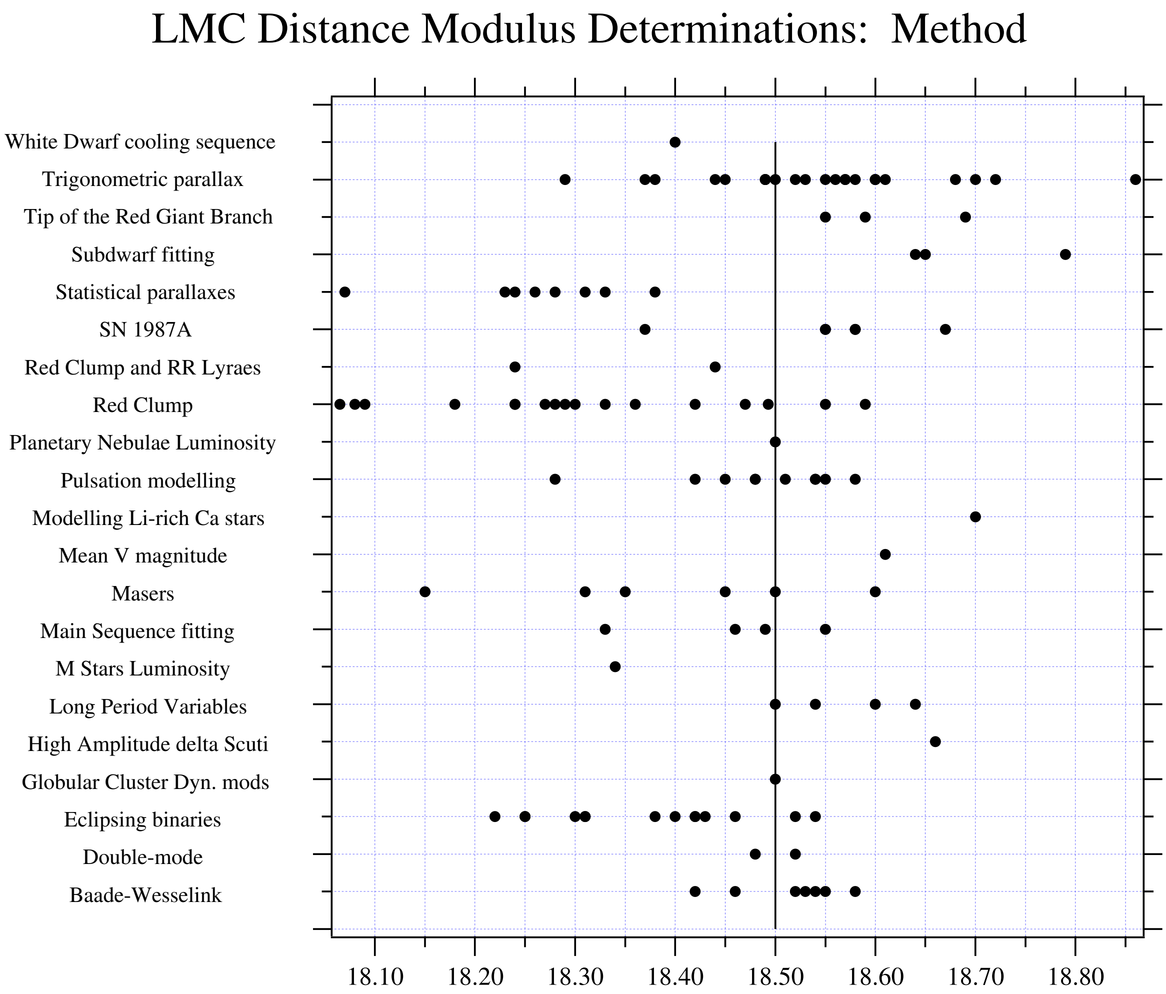

Fig. 2.— Recent determinations of the distance modulus of the Large Magellanic Cloud, sorted by method. The thick horizontal line denotes the distance modulus adopted by the HST Distance Scale Key Project (Freedman *et al.* 2001) and the Type Ia Supernovae Calibration Team (().).

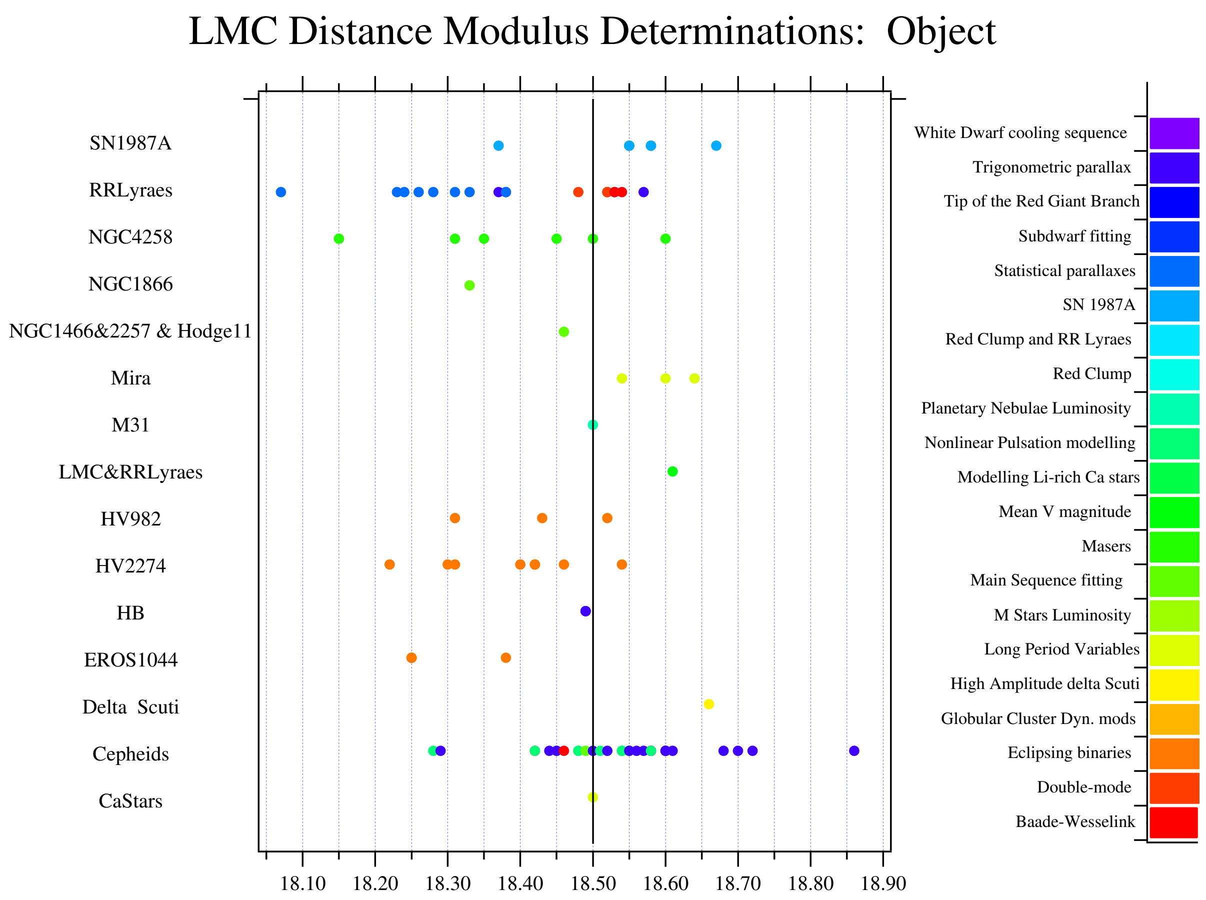

Fig. 3.— Recent determinations of the distance modulus of the Large Magellanic Cloud, sorted by object. Colors represent the various methods listed in column 2 of Table 1. The thick horizontal line denotes the distance modulus adopted by the HST Distance Scale Key Project (Freedman *et al.* 2001) and the Type Ia Supernovae Calibration Team (().).

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